A gravitationally lensed quasar with quadruple images separated by 14.62 arcseconds

Naohisa Inada¹, Masamune Oguri¹, Bartosz Pindor², Joseph F. Hennawi², Kuenley Chiu³, Wei Zheng³, Shin-Ichi Ichikawa⁴, Michael D. Gregg^{5,6}, Robert H. Becker^{5,6}, Yasushi Suto¹, Michael A. Strauss², Edwin L. Turner², Charles R. Keeton⁷, James Annis⁸, Francisco J. Castander⁹, Daniel J. Eisenstein¹⁰, Joshua A. Frieman^{7,8}, Masataka Fukugita¹¹, James E. Gunn², David E. Johnston⁷, Stephen M. Kent⁸, Robert C. Nichol¹², Gordon T. Richards², Hans-Walter Rix¹³, Erin Scott Sheldon⁷, Neta A. Bahcall², J. Brinkmann¹⁴, Željko Ivezić², Don Q. Lamb⁷, Timothy A. McKay¹⁵, Donald P. Schneider¹⁶ & Donald G. York^{7,17}

Gravitational lensing is a powerful tool for the study of the distribution of dark matter in the Universe. The cold-dark-matter model of the formation of

¹ Department of Physics, School of Science, University of Tokyo, 113-0033, Japan

² Princeton University Observatory, Peyton Hall, Princeton, NJ 08544, USA

 $^{^3}$ Department of Physics and Astronomy, Johns Hopkins University, 3701 San Martin Drive, Baltimore, MD 21218, USA

⁴ National Astronomical Observatory, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

⁵ Department of Physics, University of California at Davis, 1 Shields Avenue, Davis, CA 95616, USA

⁶ Institute of Geophysics and Planetary Physics, Lawrence Livermore National Laboratory, L-413, 7000 East Aveneu, Livermore, CA 94550, USA

⁷ Department of Astronomy and Astrophysics, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA

⁸ Fermi National Accelerator Laboratory, P.O. Box 500, Batavia, IL 60510, USA

⁹ Institut d'Estudis Espacials de Catalunya/CSIC, Gran Capita 2-4, 08034 Barcelona, Spain

 $^{^{10}}$ Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721, USA

¹¹ Institute for Cosmic Ray Research, University of Tokyo, 5-1-5 Kashiwa, Kashiwa City, Chiba 277-8582, Japan

¹² Department of Physics, Carnegie Mellon University, Pittsburgh, PA 15213, USA

¹³ Max-Planck Institute for Astronomy, Königstuhl 17, D-69117 Heidelberg, Germany

¹⁴ Apache Point Observatory, P.O. Box 59, Sunspot, NM88349, USA

 $^{^{15}}$ Department of Physics, University of Michigan, 500 East University Avenue, Ann Arbor, MI 48109, USA

¹⁶ Department of Astronomy and Astrophysics, Pennsylvania State University, 525 Davey Laboratory, University Park, PA 16802, USA

 $^{^{17}}$ Enrico Fermi Institute, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA

large-scale structures predicts¹⁻⁶ the existence of quasars gravitationally lensed by concentrations of dark matter⁷ so massive that the quasar images would be split by over 7 arcsec. Numerous searches⁸⁻¹¹ for large-separation lensed quasars have, however, been unsuccessful. All of the roughly 70 lensed quasars known¹², including the first lensed quasar discovered¹³, have smaller separations that can be explained in terms of galaxy-scale concentrations of baryonic matter. Although gravitationally lensed galaxies¹⁴ with large separations are known, quasars are more useful cosmological probes because of the simplicity of the resulting lens systems. Here we report the discovery of a lensed quasar, SDSS J1004+4112, which has a maximum separation between the components of 14.62 arcsec. Such a large separation means that the lensing object must be dominated by dark matter. Our results are fully consistent with theoretical expectations³⁻⁵ based on the cold-dark-matter model.

For bright quasars (*i*-band magnitude i < 19), the probability of gravitational lensing¹⁵ is only about 0.1%; the majority of these lenses have small separations, due to a single massive galaxy. The fraction of large-separation lensed quasars is predicted to be 0.01% or less³⁻⁵; thus it is not surprising that none have been found to date⁸⁻¹¹. In order to find such objects, we need samples of tens of thousands of quasars, such as generated by the Sloan Digital Sky Survey^{16,17} (SDSS). The SDSS is conducting both a photometric survey¹⁸⁻²² using five broad optical bands²² (u, g, r, i and z) and a spectroscopic survey²³ of 10,000 square degrees of the sky centered approximately on the North Galactic Pole, using a dedicated wide-field 2.5-m telescope at the Apache Point Observatory.

We searched for large-separation lensed quasars in a sample of $\sim 29,500$ spectroscopically-confirmed SDSS quasars²⁴ at redshifts z of 0.6-2.3, a sample roughly three times larger than those used in previous searches. Even with this large sample, the expected number of large-separation lensed quasars is of the order of unity. In the field around each quasar in the sample, we searched for stellar objects with colours differing by less than 0.1 from those of the quasar, with separations between 7".0 and 60".0 and with flux greater than one-tenth that of the quasar. SDSS J1004+4112 was identified as a 'quadruple' large-separation lensed quasar candidate using these criteria. Only one of the four components (component B, see below) has an SDSS spectrum (the SDSS hardware²³ does not allow pairs of objects separated by less than 55" to be observed on a single plate), and therefore, we obtained spectra of all four components using the Keck I telescope at the W. M. Keck Observatory. The results are shown in Fig. 1. All four components indeed show quasar-like features, with all emission lines giving a consistent redshift $z = 1.734\pm0.002$; the velocity differences of the quasar components are ~ 100 km s⁻¹, comparable to the observational uncertainty. Although it is not obvious from Fig. 1, there are C IV absorption line systems at z = 1.732 in each

of the four quasar spectra; this is an absorption system associated with the quasar itself, further supporting the lensing hypothesis: the four quasar images are from the same physical source. The differences in their spectra may be explained by the modest time-variability of the source quasar over ~ 1 year, the expected gravitational lensing time-delay²⁵ among those different images.

Additional strong support for the lensing hypothesis comes from the identification of the galaxy cluster responsible for the large-separation lensing. From the observed image separations (the maximum separation is 14''.62, between images B and C), we infer that the lensing object should have a velocity dispersion in excess of 600 km s⁻¹. Thus the lensing object cannot be a single galaxy, but must be rather a group or cluster of galaxies that has a sufficiently concentrated distribution of dark matter. To identify the lensing object, we obtained deep optical images of the system using the Subaru telescope of the National Astronomical Observatory, Japan. The result is shown in Fig. 2. A number of galaxies are clearly detected around component G, suggesting that it is the most luminous galaxy of the cluster. We obtained a spectrum of component G using the Keck I telescope. The spectrum shows a number of absorption features characteristic of a early-type galaxy at a redshift of $z = 0.6799 \pm 0.0001$. We also obtained spectra of two faint galaxies immediately to the south-west of component G (Fig. 2) using the Faint Object Camera and Spectrograph²⁶ of the Subaru telescope. The redshifts of these two faint galaxies are $z = 0.6751 \pm 0.0001$, strongly suggesting a cluster of galaxies at $z\sim0.68$ centred on component G. Clusters are dominated by elliptical galaxies, which all have very similar spectral energy distributions.

Many of the faint galaxies in Fig. 2 (\sim 40 galaxies around component G) have colours that are similar to that of component G. The colours are consistent with the expected colours of elliptical galaxies at $z\sim$ 0.68 ($g-r\sim$ 1.8 and $r-i\sim$ 1.1). In addition, there is an X-ray source in this direction detected by the ROSAT All-Sky Survey²⁷ (0.236 counts per second in a 473-s exposure). The emission, however, comes most probably from the quasar, because the detected X-ray flux is too strong for typical clusters of galaxies at z=0.68. Finally, we note two possible arclets (highly distorted images of background galaxies due to gravitational lensing) in Fig. 2 (marked as 'arc?') close to component D. If future observations confirm that the arclets are indeed lensed background galaxies, they will provide strong additional constraints on the total mass distribution of the lensing cluster.

The lensing interpretation is further supported by a theoretical model of SDSS J1004+4112. We fitted the positions of the four quasar components with a singular isothermal ellipsoid (SIE) plus external shear model using lens modeling software²⁸. The best-fit model is illustrated in Fig. 3. The positions and relative brightnesses of all components agree well with the lens model predictions. The centre of the lensing mass is offset from the centre of

component G by about 10 kpc at the cluster redshift, but brightest galaxy of a cluster is not always found exactly at the centre of the potential well of that cluster.

The identical redshifts (z=1.734) and the spectral energy distributions of the four lensed components, the existence of a lensing cluster of galaxies (z=0.68), and the presence of possible arclets confirm the hypothesis that the quasar is lensed by this cluster. Furthermore, a theoretical lensing model involving the cluster and external shear simultaneously accounts for the observed geometry of the system and the relative brightness of the images. The present work represents the discovery of a long-predicted but previously undetected population of large-separation lensed quasars.

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Correspondence and requests for materials should be addressed to N.I. (e-mail: inada@utap.phys.s.u-tokyo.ac.jp)

Table 1. ASTROMETRY AND PHOTOMETRY FOR SDSS J1004+4112

Object	R.A.(J2000) ^a	Dec.(J2000) ^a	$g^{ m b}$	$r^{ m b}$	$i^{ m b}$	$z^{ m b}$	$\Delta heta^{ m c}$
A	10 04 34.794	+41 12 39.29	18.67 ± 0.03	18.70 ± 0.02	18.46 ± 0.02	18.43 ± 0.05	373
В	10 04 34.910	$+41\ 12\ 42.79$	19.05 ± 0.06	19.10 ± 0.06	18.86 ± 0.06	18.92 ± 0.06	0'.'00
\mathbf{C}	$10\ 04\ 33.823$	$+41\ 12\ 34.82$	19.71 ± 0.03	19.73 ± 0.02	19.36 ± 0.03	$19.31 {\pm} 0.07$	14.62
D	$10\ 04\ 34.056$	$+41\ 12\ 48.95$	20.67 ± 0.04	20.51 ± 0.04	20.05 ± 0.04	20.00 ± 0.13	1144
G	$10\ 04\ 34.170$	$+41\ 12\ 43.66$	22.11 ± 0.40	$20.51 {\pm} 0.13$	19.54 ± 0.09	19.04 ± 0.21	8.44

^aR.A., right ascension; Dec., declination. RA and Dec. are given in the J2000 system. These celestial coordinates were measured on the basis of the celestial coordinates of component B. The positional errors of components A, C, and D (not including the absolute positional errors of component B) are 0.001 and that of component G is 0.002 per coordinate.

 $^{^{\}mathrm{b}}g$, r, i and z mean the magnitudes of each band.

^cSeparation angles relative to component B

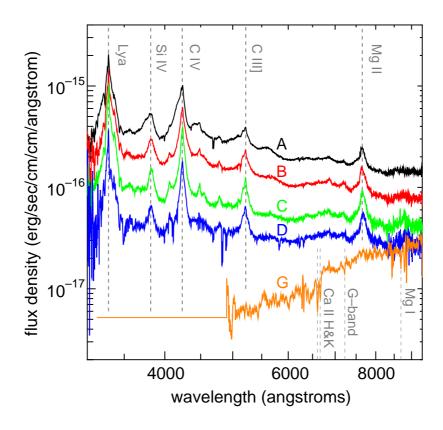


Figure 1 The Keck spectra of the four quasar components A–D and the brightest galaxy G in the lensing cluster. See Fig. 2 for these identifications (A–D, and G). The data were taken using the Low-Resolution Imaging Spectrometer²⁹ (LRIS) of the Keck I telescope. The exposure times were 900 s for each component. The dispersion is 1.09Å pixel⁻¹. The data were reduced in a standard method using IRAF (IRAF is the image reduction and analysis facility, distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation). The black solid line, the red solid line, the green solid line, and the blue solid line represent the spectra of components A, B, C, and D, respectively. The vertical gray dotted lines (3323.6 Å, 3818.7 Å, 4235.1 Å, 5218.5 Å, and 7651.8Å) represent the positions of emission lines of the respective ions redshifted to z = 1.734 of Ly α (1215.67 Å), Si IV (1396.76 Å), C IV (1549.06 Å), C III] (1908.73 Å), and Mg II (2798.75 Å), respectively. All emission lines are clearly at the same redshift. The orange solid line represents the Keck spectrum of component G at the same dispersion. The exposure time was also 900 s for component G. The vertical thinner gray dotted lines (3933.7 Å, 3968.5 Å, 4304.4 Å, and 5175.3 Å) represent the positions of absorption lines of the respective ions redshifted to z = 0.680 of Ca II H&K (3933.7 Å and 3968.5 Å), G-band (4304.4 Å), and Mg I b-band (5175.3 Å), respectively. There are no data below ~ 4900 Å in the spectrum of component G.

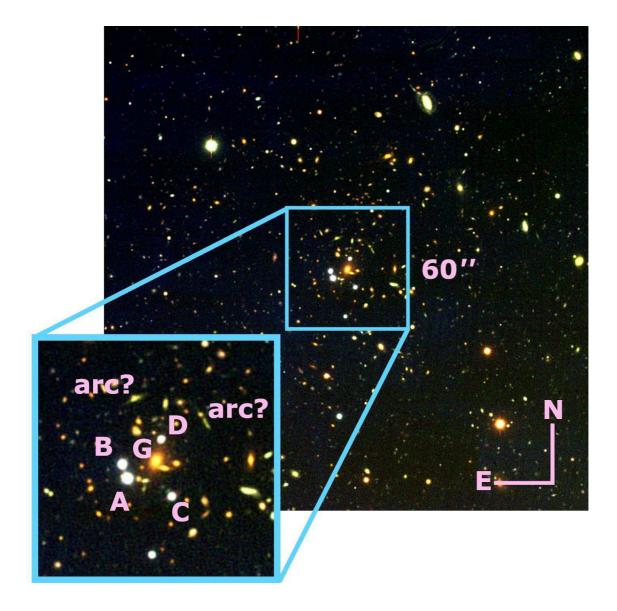


Figure 2 The gri composite Subaru image of the field around SDSS J1004+4112. The data were taken using the Subaru Prime Focus Camera³⁰ of the Subaru telescope. The magnitude limit is $i \approx 26.0$. The central 60" square is shown in an expanded view. The four quasar components are marked as A, B, C and D, and the bright galaxy located between the four quasar components is marked as G. The separation between components A and D is 12".77, and that between components B and C is 14".62. The positions (J2000) and the magnitudes of the components A–D and the brightest galaxy (component G) between the four quasar components are summarized in Table 1. Many faint galaxies can be seen — their positions and colours are consistent with being members of a cluster (z = 0.68) centred on component G. Two possible arclets (marked as 'arc?') can also be seen. The seeing had a full-width at half-maximum of 0".6.

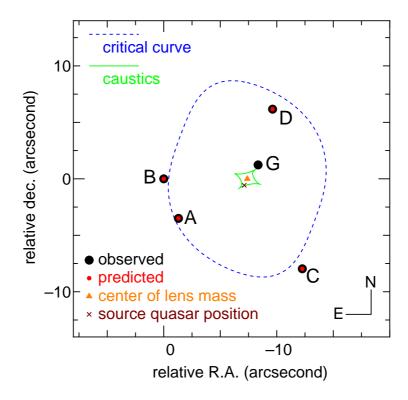


Figure 3 The best fit lens model prediction compared with the observation. We used a lensing model of a singular isothermal ellipsoid (SIE) with external shear. The best-fit model has an Einstein radius of the SIE model $\alpha_e = 6^{\prime\prime}.906$ (corresponding to a velocity dispersion ~ 700 km sec⁻¹ at the cluster redshift), magnitude and position angle of the shear $\gamma = 0.250$ and $\theta_{\gamma} = -60.925^{\circ}$ (measured East of North), and ellipticity and its position angle e = 0.498 and $\theta_e = 21.434^{\circ}$ (measured East of North), with a source quasar position ($\Delta R.A.$, $\Delta Dec.$)=(-7".124, -0".574) and a centre of lensing mass $(\Delta R.A., \Delta Dec.) = (-7''.387, -0''.004)$ relative to the centre of component A. The black filled circles represent the observed positions of components A, B, C, D and G, and the red filled circles represent the predicted positions of components A–D. The green solid line is the position of the caustic in the source plane, and the blue dashed line represent the critical curve in the image plane. The small brown filled cross is the predicted position of the source quasar, and the small orange triangle is the predicted position of the center of the lens mass. The differences between the observed and modelled image positions are much smaller than the observational uncertainties. The flux ratios predicted from the model, B/A, C/A and D/A are 0.78, 0.43 and 0.22, respectively. The total magnification of the quasar images which is predicted by the model is 56.48. The predicted flux ratios are close to the observational results; $B/A = 0.69 \pm 0.04$, $C/A = 0.46 \pm 0.02$, and $D/A = 0.25 \pm 0.01$ (measured from the i band image). Microlensing by substructures and/or reddening by the Mg II absorption line systems that are seen in each spectrum might be the cause of the differences between the predicted flux ratios and the observations.